Sgr A*の過去の活動を起源とする銀河系中心からのガンマ線放射

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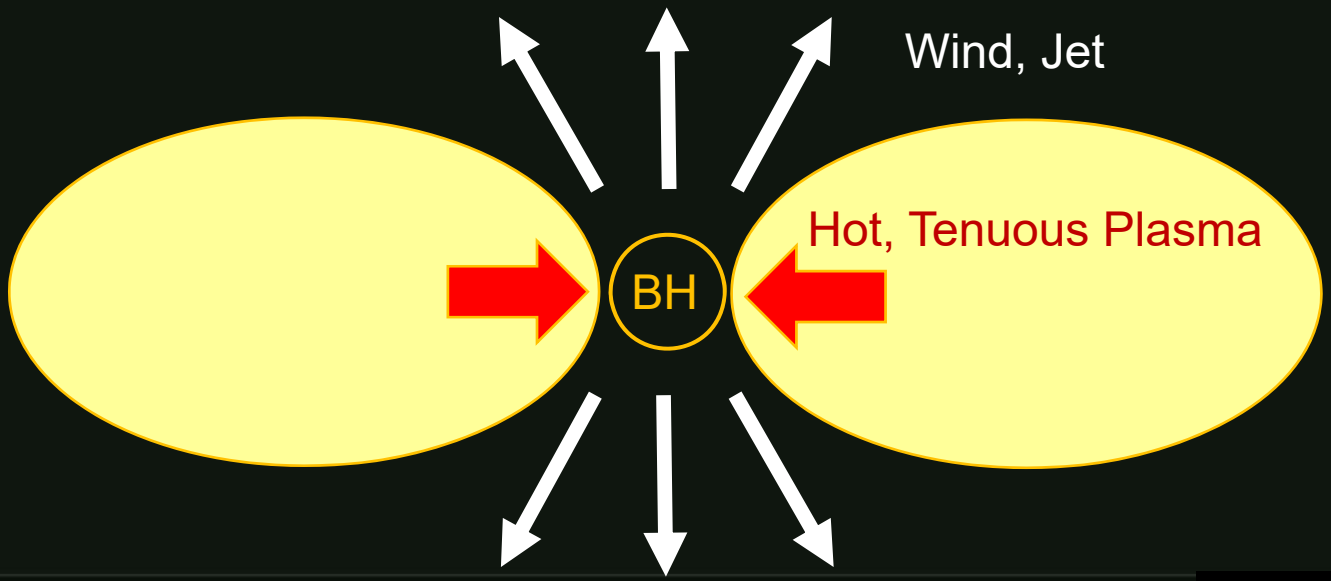
- Summary

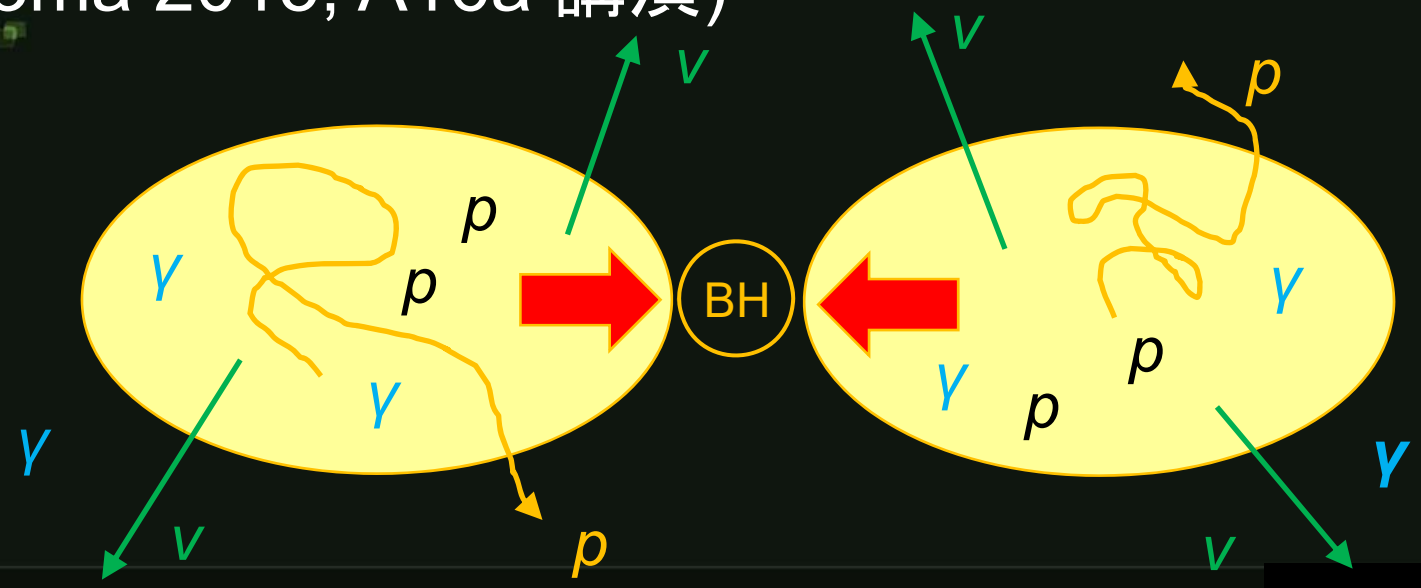
CR acceleration in RIAF

- Accretion flows for LLAGN are thought to be RIAFs
  - Hot and tenuous
  - Coulomb collision is inefficient (collisionless)
    - Insufficient thermalization → particle acceleration



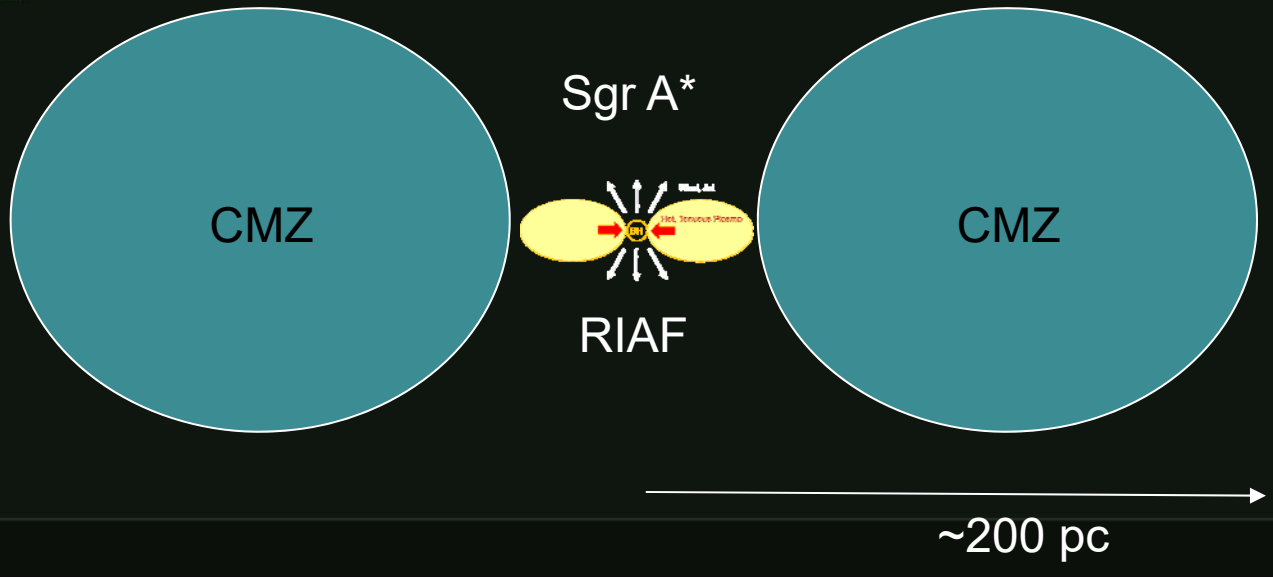
CR acceleration in RIAF

- Particles are stochastically accelerated by turbulence in RIAF
  - *pγ* and *pp*-interactions in the RIAFs may be responsible for production of the neutrinos detected by IceCube (Kimura, Murase, & Toma 2015; A16a 講演)



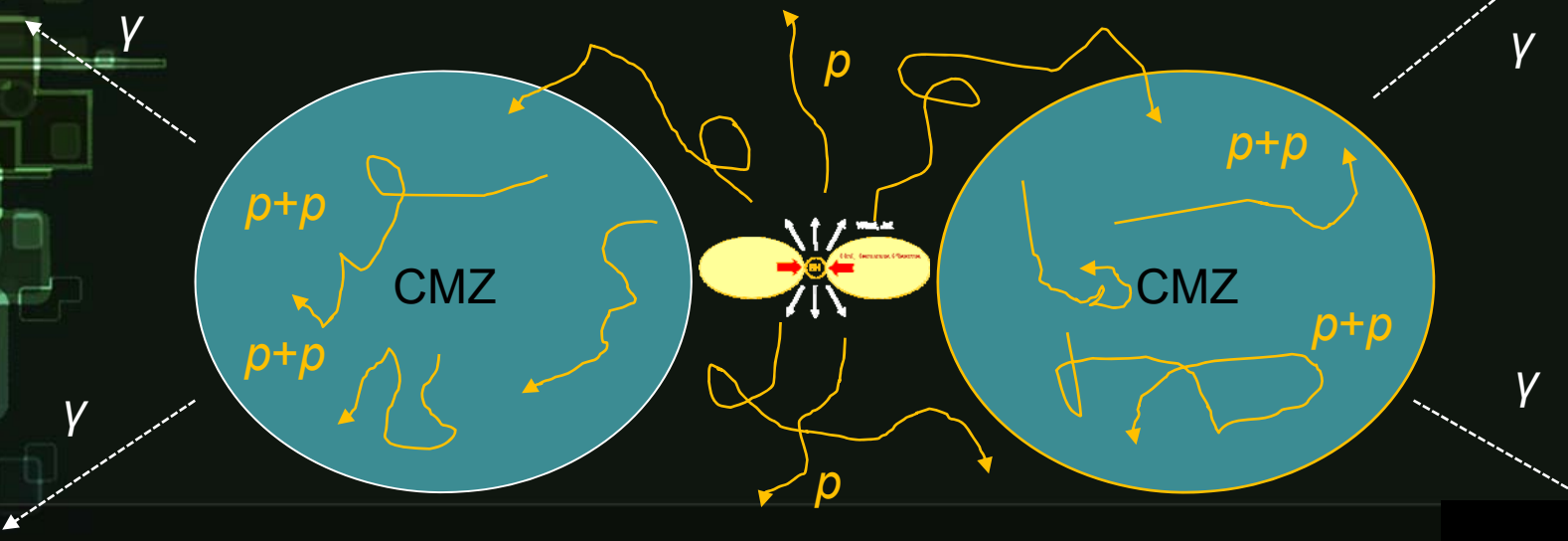
Sgr A* and CMZ

- Sgr A* is the supermassive black hole at the Galactic center
  - It is a LLAGN and has a RIAF
    - CRs may be accelerated
  - Surrounded by molecular gas (CMZ)
    - Mass of CMZ is ~ 10^7 $M_\odot$



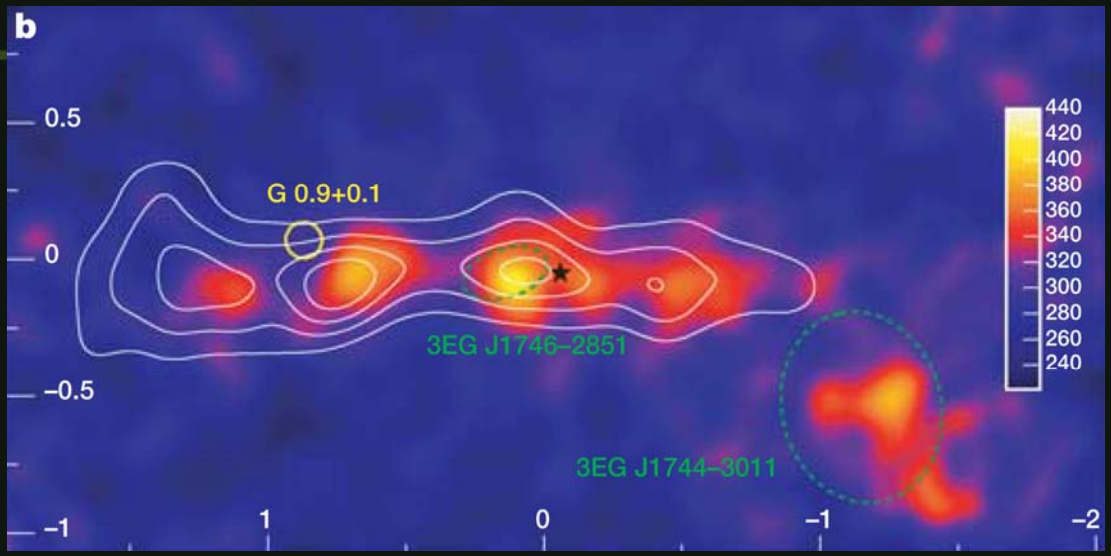
CR protons in CMZ

- Some of the protons accelerated in the RIAF should plunge into CMZ
  - They produce γ-rays through *pp*-interactions
  - We calculate the diffusion of CR protons in the CMZ and the γ-ray emission



γ-rays from the CMZ

- HESS observation of the Galactic center
  - CMZ radiates γ-rays



Color: γ-ray
Contour: molecular gas
(Aharonian+ 06)

Model

- We solve diffusion equation
  - Spherically symmetric
    - We ignore the region perpendicular to the disc

$$\frac{\partial f}{\partial t} = \frac{1}{r^2}\frac{\partial}{\partial r}\left(r^2 \kappa \frac{\partial f}{\partial r}\right) + Q$$

  - Diffusion coefficient (κ)
    - Typical Galactic value (Gabici+ 09)

$$\kappa = 10^{28}\left(\frac{E_p}{10\ \mathrm{GeV}}\right)^{0.5}\left(\frac{B}{3\ \mu\mathrm{G}}\right)^{-0.5}\ \mathrm{cm^2 s^{-1}}\ ,$$

    - B ~ mG in CMZ
  - Source (Q)
    - CRs accelerated in Sgr A* (Kimura+ 13)
    - Only a tiny fraction of CRs plunge into the CMZ

      $\lambda \sim 10^{-3}$

Spectrum of CRs

- The typical energy of CRs accelerated in a RIAF is determined by the condition of $t_{\rm acc} = t_{\rm diff,RIAF}$

$$\frac{E_{p,\rm eq}}{m_p c^2} \sim 1.4 \times 10^5 \left(\frac{\dot{m}}{0.01}\right)^{1/2} \left(\frac{M_{\rm BH}}{1 \times 10^7\, M_\odot}\right)^{1/2} \left(\frac{\alpha}{0.1}\right)^{1/2} \left(\frac{\zeta}{0.1}\right)^{3} \left(\frac{\beta}{3}\right)^{-2} \left(\frac{R_{\rm acc}}{10\, R_S}\right)^{-7/4}$$

- Functional form
  - 2nd Fermi acceleration (Becker+ 06)

$$\dot{N}(x)dx \propto x^{(7-3q)/2} K_{(b-1)/2}(x)dx$$

K_ν: Bessel func.

- Normalization

$$\propto \dot{M}c^2$$

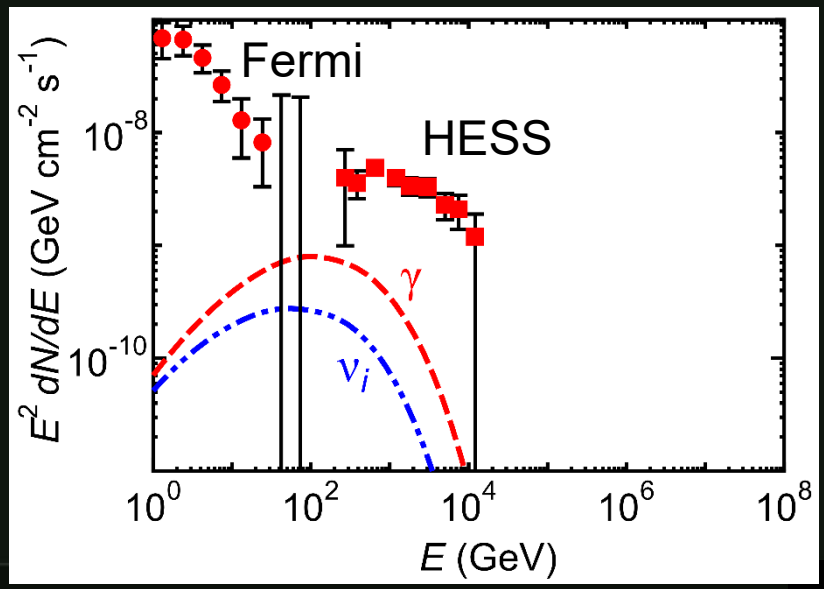
Results

- Current accretion rate on Sgr A* is very small

  $\dot{m} = \dot{M}/\dot{M}_{\rm Edd} = 4.2 \times 10^{-6}$

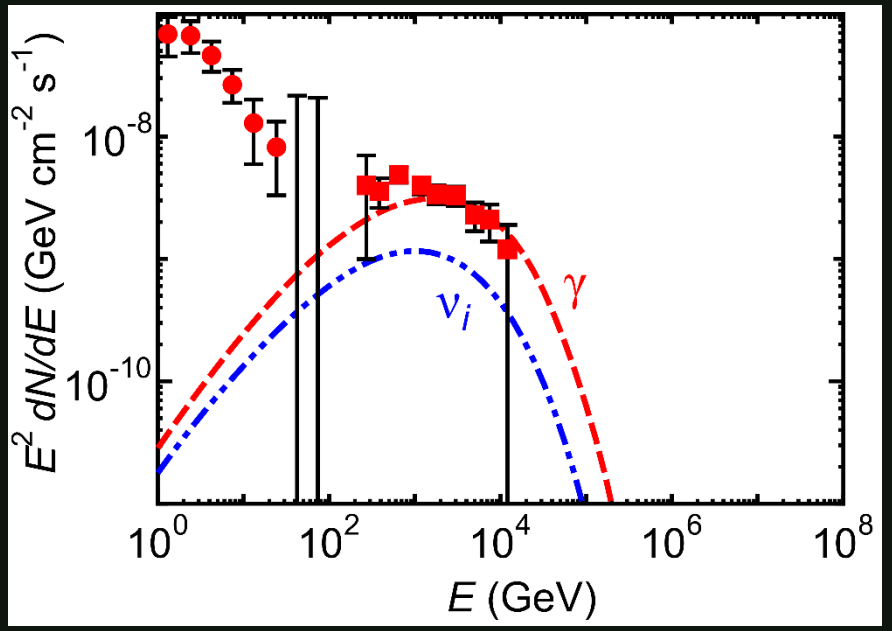
- If $\dot{m}$ is constant, γ-ray luminosity is much smaller than the Fermi and HESS observations

Spectrum
Red: γ-ray
Blue: neutrino



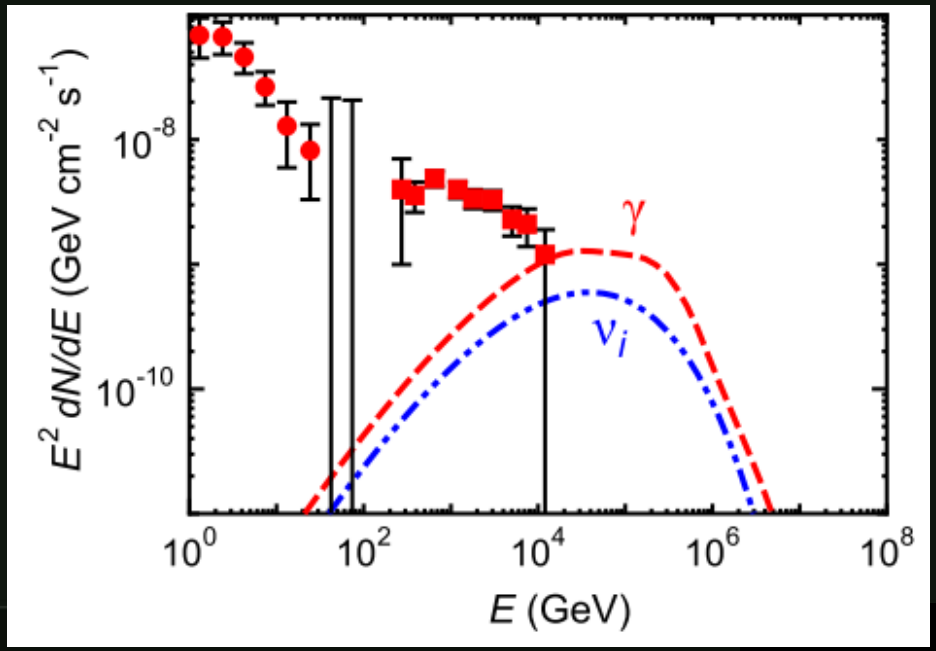
Past activity of Sgr A*

- Observations have indicated that Sgr A* was much more active in the past
  - Koyama+96, Murakami+00, Totani+06, Ryu+13
    - $\dot{m} \sim 0.001$ ($\gtrsim$ 100 yrs ago)
- TeV γ-ray observations can be explained
  - GeV emission has another origin



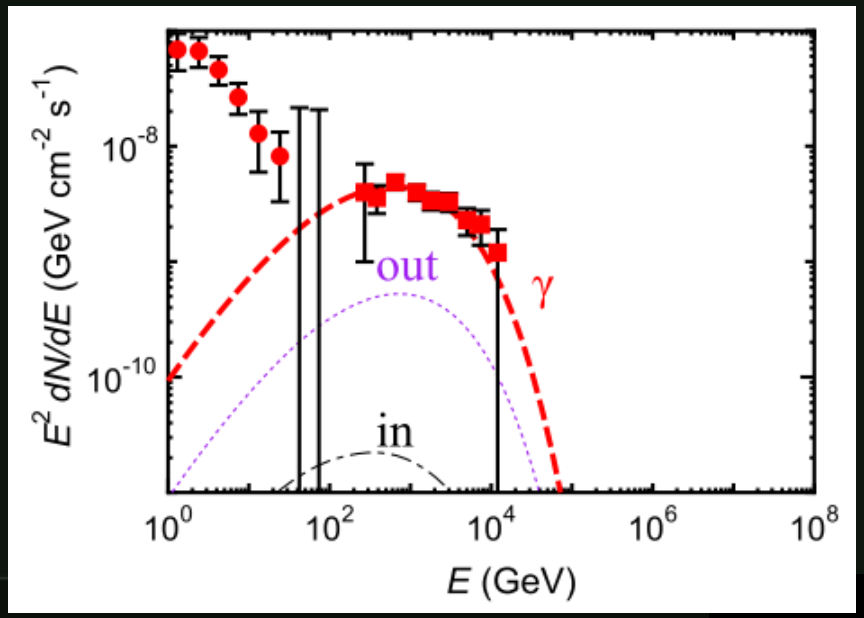
Uncertainty

- Since our model is rather simple, there are uncertainties in parameters
  - Increase the typical energy of protons assuming more efficient acceleration
  - CTA can discriminate the models



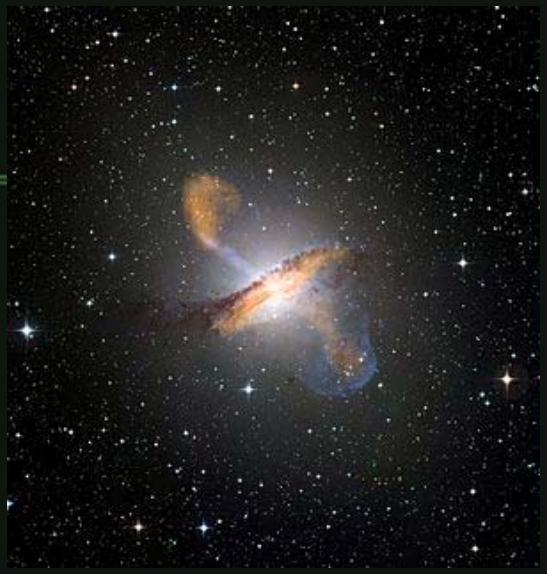
History of CR Acceleration

- In our model, γ-rays from the outer CMZ are radiated by older CRs
  - It takes a longer time for the CRs to diffuse
- Spectral variance of γ-ray spectrum may reflect the history of CR acceleration in Sgr A*
  - Change of typical energy of CRs

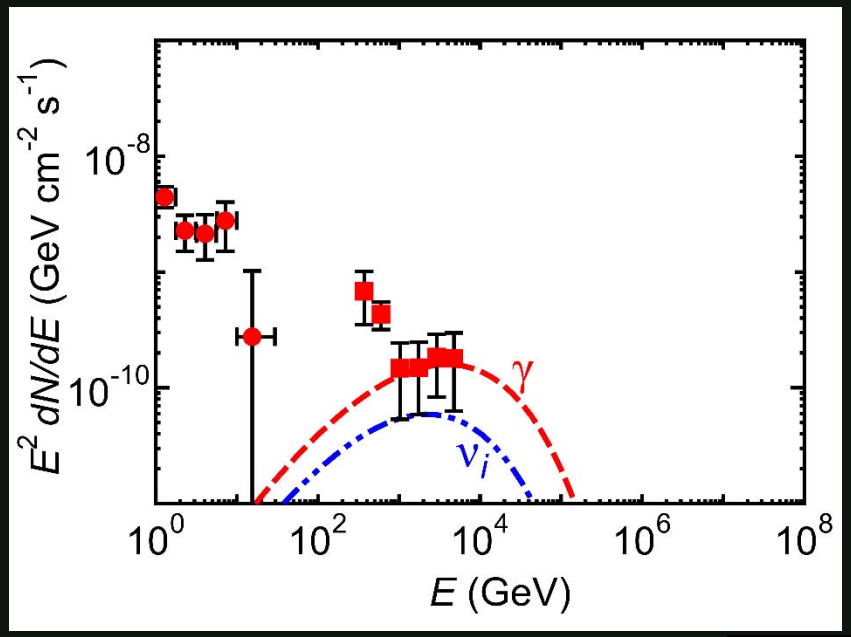


Other galaxies

- Centaurus A
  - Nearby radio galaxy (3.84 Mpc)
  - TeV γ-ray emission can be explained if the AGN is well covered by molecular gas
    - $\lambda \sim 0.02$



(by ESO)



Summary

- LLAGNs may accelerate CRs and may be the source of neutrinos detected by IceCube
  - If so, Sgr A* should produce a lot of CR protons
  - Some of them should enter the CMZ and generate γ-rays
- We solved diffusion equation of the CRs and calculated γ-ray spectra
  - The results are consistent with TeV γ-ray observations if Sgr A* was active in the past